



## SCIENCE RESULTS

# UVIT study of UV bright stars in the globular cluster NGC 4147

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**Abstract.** We present far ultraviolet (FUV) observations of globular cluster NGC 4147 using three FUV filters, BaF<sub>2</sub> (F154W), sapphire (F169M), and silica (F172M) of Ultra-Violet Imaging Telescope (UVIT) on-board the AstroSat satellite. We confirmed the cluster membership of the UVIT observed sources using proper motions from Gaia data release 2 (GAIA DR2). We identified 37 blue horizontal branch stars (BHBs), one blue straggler star (BSS) and 15 variable stars using UV-optical color magnitude diagrams (CMDs). We find that all the FUV bright BHBs are second generation population stars. Using UV-optical CMDs, we identify two sub-populations, BHB1 and BHB2, among the UV-bright BHBs in the cluster with stars count ratio of 24:13 for BHB1 and BHB2. The effective temperatures ( $T_{\text{eff}}$ ) of BHB1 and BHB2 were derived using color-temperature relation of BaSTI-IAC zero-age horizontal branch (ZAHB). We found that BHB1 stars are more centrally concentrated than BHB2 stars. We also derive physical parameters of the detected FUV bright BSS by fitting younger age BaSTI-IAC isochrones on optical and UV-optical CMDs.

**Keywords.** Galaxy: globular clusters: individual: NGC 4147—stars: horizontal-branch—stars: blue stragglers—stars: Hertzsprung–Russell and colour-magnitude diagrams.

## 1. Introduction

The glimpse of ultraviolet (UV) light in the old stellar population of the Galactic globular clusters (GGCs) is dominated by hot luminous UV-bright stars which are mostly the stars of blue horizontal branch (BHB), blue-stragglers (BS), post-asymptotic giant branch (pAGB), and extreme horizontal branch (EHB) phases having temperature more than 7000 K. Various physical properties of these UV-bright populations have been explored using large sample of GGCs observed by *Galaxy Evolution Explorer* (GALEX, Schiavon *et al.* 2012) and *Hubble space telescope*

(*HST*, Nardiello *et al.* 2018). Ultraviolet Imaging Telescope (UVIT) on-board the AstroSat satellite (Kumar *et al.* 2012) has also performed the imaging observations of several GGCs in UV with a better spatial resolution than *GALEX* enabling in resolving the core of the clusters and also in distinguishing UV bright stars of different evolutionary stages to study their physical parameters in a great detail (Subramaniam *et al.* 2017; Sahu *et al.* 2019a; Jain *et al.* 2019; Kumar *et al.* 2020, 2021; Rani *et al.* 2020; Singh *et al.* 2020). Using the spectroscopic data of a large number of stars in many GGCs along with the UV photometric observations with *HST*, the evidence for the presence of multiple stellar populations in GGCs is now well established (see reviews Bastian & Lardo 2018; Gratton *et al.* 2019; Cassisi & Salaris 2020, and

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references therein). The different chemical evolution within the GGCs seems to be the origin of two distinct sub-population of stars (the first generation (1G) and second generation (2G)) characterized by significant abundance variations in light elements. The stellar populations enriched in He, N, and depleted in O and C are 2G stellar population while stellar populations enriched in C and O and depleted in N and Na with a primordial He-abundance are termed as normal or 1G stellar population (see Milone *et al.* 2017; Marino *et al.* 2019, and references therein).

We present here the UV study of an old age, metal poor, low density GC NGC 4147, located at a distance of 21 kpc from the Galactic center and 19 kpc from the Sun ( $l = 252.84^\circ$ ,  $b = +77.18^\circ$ ). We have observed this cluster using three far-UV (FUV) filters of UVIT. Prior to this, the UV imaging and color-magnitude diagrams (CMDs) of the cluster were presented by Schiavon *et al.* (2012) using *GALEX* observations. Their study was restricted to the sources in the outer region of the cluster due to the lower resolution ( $\sim 5''$ ) of *GALEX*. However, the cluster is well studied both photometrically and spectroscopically in the optical bands of the electromagnetic spectrum. Several of the photometric observations in the optical bands have explored about specific sources of the cluster such as BHBs, variable stars, red giant branch stars (RGBs), etc., using their optical CMDs (Auriere & Lauzeral 1991; Arellano Ferro *et al.* 2004; Stetson *et al.* 2005; Arellano Ferro *et al.* 2018; Lata *et al.* 2019). Similarly, many low resolution spectroscopic observations of the cluster have been performed to find its metallicity ( $[\text{Fe}/\text{H}] = -1.85$  dex), and chemical abundances of various alpha-elements (Suntzeff *et al.* 1988; Martell *et al.* 2008; Ivans 2009). The latest high-resolution spectroscopy of 18 RGB stars by Villanova *et al.* (2016) has revealed that most of the RGB stars of the cluster are of 2G population type with 2G to 1G RGB stars' ratio of 85:15. They found that the red-HB stars (RHBs) are progeny of the 1G population and the BHBs are progeny of the 2G population. They also confirmed an alpha-enhancement of 0.38 dex and a mono-metallicity of  $[\text{Fe}/\text{H}] = -1.84$  dex among 1G and 2G stars in the cluster.

In Section 2, we present the observation details, data reduction process and photometry of the UVIT observation of the cluster. In Section 3, we show the CMDs of all the detected sources. In Section 4, we discuss the sub-population among the UVIT detected 2G BHBs. In Section 5, we estimate various physical properties of the UV bright blue straggler star (BSS). Finally, we summarize our results in Section 6.

## 2. Data reduction

We obtained the UVIT observation of the cluster NGC 4147 in three FUV filters, BaF<sub>2</sub>, sapphire and silica, having wide, medium and narrow bandwidths, respectively. The UVIT data of the cluster was obtained from AstroSat archival web-page<sup>1</sup> in Level 1 format. The instrument and calibration details of UVIT can be found in Tandon *et al.* (2017, 2020). We reduced the Level 1 data into science image using a customized software package CCDLAB (Postma & Leahy 2017), specially designed for the UVIT data reduction. The UVIT observations have been performed in several orbits and we get science image for each orbit. In order to get a good signal to noise ratio (SNR), we combined all the observed orbits into a single image and then performed the photometry. The UVIT observation details are given in Table 1.

In Fig. 1, we provide a color image of the cluster using multi-wavelength observations from infrared (IR) to UV. The sources in blue color are observed in UV using the UVIT BaF<sub>2</sub> filter, the green color sources are from the archival catalog of CFHT-4m *V* band observation (Stetson *et al.* 2019), and the red color sources are obtained from 2MASS *J* band observation<sup>2</sup> of the cluster. We can see that the hot sources, observed in UV (blue sources) are bright enough and easily distinguishable in the outer region of the cluster. The central region of the cluster also contains many hot sources as visible in cyan color (mixture of blue and green) due to crowding effect at the center.

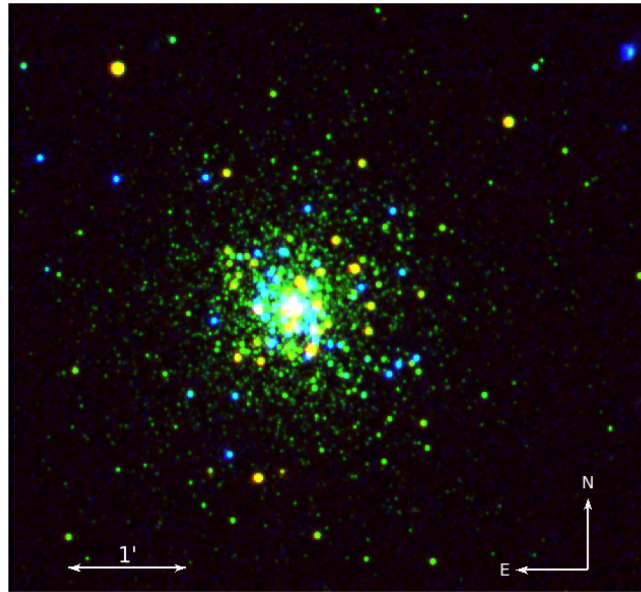
We applied DAOPHOT point spread function (PSF) photometry (Stetson 1987) on the science images to obtain the source positions and their respective apparent magnitudes. The aperture and saturation corrections were applied on the detected sources in all the three FUV filters following the suggestion by Tandon *et al.* (2017). A typical PSF of  $1.5''$  was obtained in all the science images of the FUV filters. We have excluded the sources within  $10''$  from the cluster center to avoid contamination. We detected 114, 92, and 65 sources in the BaF<sub>2</sub>, sapphire, and silica filters, respectively. The AB-magnitude limits of the detected sources in BaF<sub>2</sub>, sapphire and silica filters are 23.0, 22.5, and 21.0, respectively. The magnitudes were corrected for extinction value using  $E(B - V) = 0.0221$  mag (Schlafly & Finkbeiner 2011) along the

<sup>1</sup>[https://astrobrowse.issdc.gov.in/astro\\_archive/archive/Home.jsp](https://astrobrowse.issdc.gov.in/astro_archive/archive/Home.jsp).

<sup>2</sup><https://irsa.ipac.caltech.edu/applications/2MASS>.

**Table 1.** Observational details of NGC 4147.

Date of observation: 2017, April 17					
Telescope pointing: RA = 182.5263°, DEC = 18.5426°					
Channel	Filters	Mean $\lambda$ (Å)	$\Delta\lambda$ (Å)	Exp. time (s)	No. of orbits
FUV	BaF <sub>2</sub>	1541	380	1536	5
	Sapphire	1608	290	1648	3
	Silica	1717	125	1209	2



**Figure 1.** Color image of NGC 4147; UV: UVIT BaF<sub>2</sub> (blue), visible: CFHT-4m V band (green), infrared: 2MASS J band (red).

cluster direction and extinction law of Cardelli *et al.* (1989) in the UV filters.

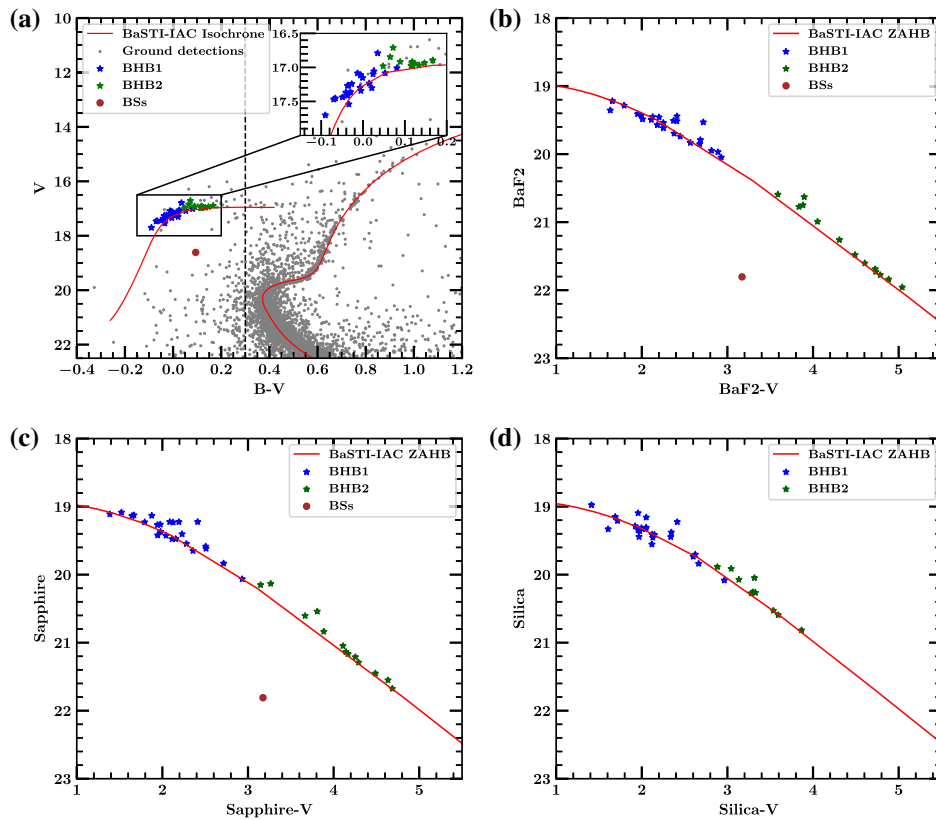
### 3. Color-magnitude diagrams

To study UV and UV-optical CMDs of the cluster members, we cross-matched the UVIT observed sources with GAIA data release 2 catalog (Gaia DR2, Gaia Collaboration *et al.* 2018) with a matching radius of 1.5". The confirmed cluster members were separated out from the field stars and the background sources (galaxies, quasars, etc.) using Gaia DR2 proper motions. Once the sources were confirmed as cluster members, we cross-matched them to deeper photometric observations, the latest released archival catalog of ground-based observations in UBVRI filters (Stetson *et al.* 2019). An optical CMD,  $B-V$  vs.  $V$ , is

shown in panel (a) of Fig. 2 for all the ground-based observed sources within the UVIT field of view (FoV  $\sim 30'$ ) in gray solids. The UVIT detections are over-plotted in the CMD and are shown in blue and green asterisks and in brown solid circle.

In order to obtain the cluster properties, we used the Bag of Stellar Tracks and Isochrones models (BaSTI-IAC<sup>3</sup>, Hidalgo *et al.* 2018) and generated isochrones and zero-age horizontal branch (ZAHB) with cluster parameters,  $[Fe/H] = -1.896$ , age = 13.0 Gyr (Harris 2010) and  $[\alpha/Fe] = 0.40$  (Villanova *et al.* 2016). Since the alpha-enhanced set of the BaSTI-IAC library has not been still published, we relied on the extension of the BaSTI-IAC library to alpha-enhanced mixture (Pietrinferni *et al.* 2021) for the present analysis. In panel (a) of Fig. 2, BaSTI-IAC isochrones

<sup>3</sup><http://basti-iac.oa-abruzzo.inaf.it/index.html>.



**Figure 2.** Optical and UV-optical CMDs of the UVIT observed sources of NGC 4147. *Panel (a):* The  $B-V$  vs.  $V$  CMD where ground-based detections from Stetson *et al.* (2019) are shown in gray solids, the UVIT observed BHBs are shown in blue and green asterisks. The BSS is shown in the brown solid. BaSTI-IAC isochrones and ZAHB with cluster parameter,  $[\text{Fe}/\text{H}] = -1.896$ ,  $[\alpha/\text{Fe}] = 0.40$ , Age = 13.0 Gyr, and distance-modulus ( $m - M$ ) = 16.40, are over-plotted on the observed sources. *Panels (b), (c) and (d):*  $\text{BaF}_2-V$  vs.  $\text{BaF}_2$ , sapphire- $V$  vs. sapphire, and silica- $V$  vs. silica CMDs, respectively, are shown for the UVIT observed sources with the same colors as mentioned in panel (a), and BaSTI-IAC ZAHB with the same cluster parameters as mentioned in panel (a) are over-plotted on all CMDs.

are over-plotted on the  $B-V$  vs.  $V$  CMD in red line using a distance-modulus of 16.40, which are fitting very well with the ground-based observations. In panels (b), (c), and (d) of Fig. 2, we have shown the UV-optical CMDs of the cluster using  $\text{BaF}_2$ , sapphire, and silica filters, respectively, in combination with  $V$  magnitude from the optical bands. The ZAHB from BaSTI-IAC library is over plotted on all the UV-optical CMDs. Although all the UV-optical CMDs look similar we have shown them to see the morphological distribution of UV bright stars of the cluster in all the FUV filters. We can see in Fig. 2 that BHBs are major contributors to the UV bright sources of the cluster.

If we look at the optical CMD in panel (a) of Fig. 2, we notice that the UVIT observed BHBs have spread along the horizontal line of the HB phase. This spread is better visible in UV-optical CMDs along the diagonal HB line than the optical CMD. In fact, the spread among BHBs in  $\text{BaF}_2-V$  vs.  $\text{BaF}_2$  CMD (panel (b)) becomes more clear with a gap of 0.8 magnitude in the

$\text{BaF}_2-V$  color within BHBs. Based upon the gap observed in  $\text{BaF}_2-V$  vs.  $\text{BaF}_2$  CMD, we divide the UVIT observed BHBs in two groups: BHBs with  $\text{BaF}_2-V$  color brighter than 2.8 magnitude as BHB1 (blue asterisks) and BHBs with  $\text{BaF}_2-V$  color fainter than 3.6 magnitude as BHB2 (green asterisks). These BHBs are lying in left and right portions of the BHB region in  $B-V$  vs.  $V$  CMD (panel (a)) and they are clearly identified in sapphire- $V$  vs. sapphire and silica- $V$  vs. silica CMDs in panels (c) and (d), respectively. There are 24 sources in BHB1 and 13 sources in BHB2 groups.

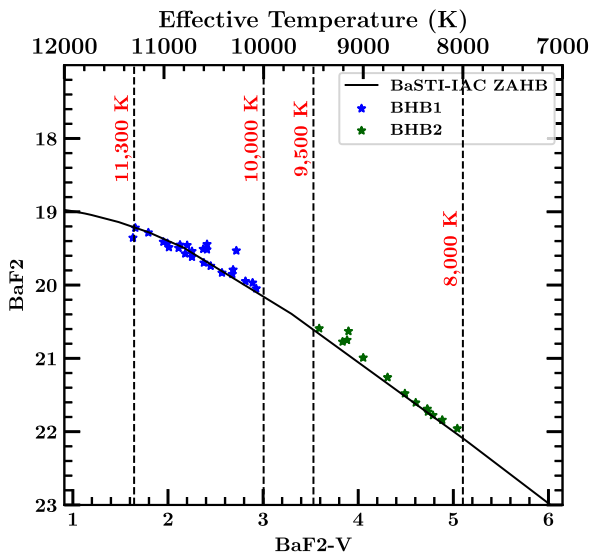
Villanova *et al.* (2016) have made a chemical abundance analysis of 18 RGB stars of the cluster NGC 4147 and found that the cluster is composed of 1G stars with  $[\text{Na}/\text{Fe}] \sim +0.0$  and  $[\text{O}/\text{Fe}] \sim +0.3$  and 2G stars with  $[\text{Na}/\text{Fe}] \sim +0.5$  and  $[\text{O}/\text{Fe}] \sim -0.2$ . In Fig. 1 of their paper they have shown that the 2G HB stars are located at  $B-V < 0.3$ . If we compare the UVIT observations with optical

data, it appears that all the observed BHBs seem to belong to the 2G population (panel (a) of Fig. 2). In the same figure, one can notice that the BHBs are separated in two sub-groups, which could point to the presence of two sub-populations in the observed 2G BHBs. We discuss this possibility in the following sections.

Apart from the BHBs, we have also detected one UV bright BSS in the cluster. The BSS was observed in BaF<sub>2</sub> and sapphire filters but not in the silica filter. The brightness of the BSS may be fainter than the detection limit of silica filter (21.0 AB-magnitude). The BSS is shown in Fig. 2 with brown solid. We also found 15 variable stars by cross-matching the UVIT observed sources with the updated variable stars catalog of GCs (Clement 2017). However, we have not shown these variable stars on the CMDs due to the severe crowding conditions affecting these stars, and the fact that we do not have a complete sampling of their light curves.

#### 4. Sub-population in BHBs

We derive the effective temperature ( $T_{\text{eff}}$ ) of the observed BHBs using color-temperature relation of the BaSTI-IAC ZAHB in Fig. 3. The best fitted ZAHB from Fig. 2 is taken to derive the color-temperature relation of the UVIT observed BHBs. As shown in



**Figure 3.** BaF<sub>2</sub>-V vs. BaF<sub>2</sub> CMD of the observed BHBs. We use BaSTI-IAC ZAHB color-temperature relation to derive the range of  $T_{\text{eff}}$  of BHB1 and BHB2. The  $T_{\text{eff}}$  derived from the ZAHB is given in the upper  $x$ -axis of the plot.

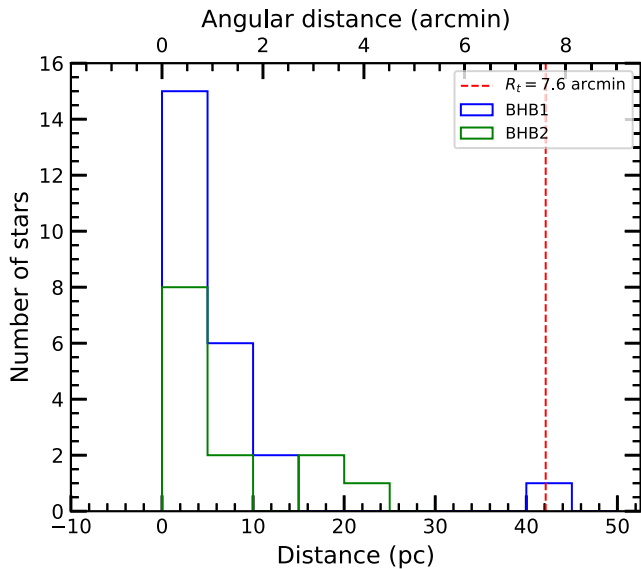
Fig. 3, the  $T_{\text{eff}}$  of BHB1 is ranging between 10000 K and 11300 K and  $T_{\text{eff}}$  of BHB2 ranges between 8000 K and 9500 K. The comparison between observations and models suggests the presence of a gap between BHB1 and BHB2 stars with a temperature difference of about 500 K. However, we caution that this width of the observed temperature gap and its actual existence may be affected by the limited number of sample stars (number of BHBs detected) in this cluster. Although a detailed statistical analysis of the reality of such a gap will be carried out in a future work, we note that a puzzling gap or discontinuity in the distribution of BHBs is seen in the color-magnitude diagram (CMD) of many GGCs (Catelan *et al.* 1998; Piotto *et al.* 1999; Behr *et al.* 2000; Brown *et al.* 2016). It was initially believed to be due to statistically significant under-population of stars but its appearance at similar location for different clusters has substantiated the claim that it is indeed a real feature and the satisfactory explanation for its origin is not yet known. However, speculation is that the gaps may be demarcating the boundaries between separate, discrete populations of HB stars, which differ in their origin or evolution.

By relying on the BaF<sub>2</sub>-V vs. BaF<sub>2</sub> CMDs obtained in the present work, we collect some evidence suggesting the possible presence of a bi-modality among the BHB population. In the following, we provide some suggestions about the possible origin for the observed discontinuity in the color distribution of BHB stars in NGC 4147. The main empirical findings are:

- (i) the BHB1 and BHB2 are clustered in the optical plane in two quite clear separate regions – although in that plane these two regions appear quite adjacent (panel (a) of Fig. 2);
- (ii) in the UV-optical plane there exists an evident gap of about 500 K whose width has been evaluated by over-imposing the models.

These empirical evidence suggests that BHB1 stars could be slightly He enhanced with respect to BHB2 ones and in such a case they would represent two distinct sub-populations in the 2G population; or alternatively they could share with BHB1 more or less the same He abundance, and have experienced different mass loss efficiency during the previous RGB stage (see, e.g., Tailo *et al.* 2020).

In Fig. 4, we have shown the radial distribution of BHB1 and BHB2 stars. We see that the BHB1 stars are spread within 15 pc (with an exception of one BHB1 star at a distance of 43 pc) from the cluster



**Figure 4.** The radial distribution of UVIT detected BHBs. The histogram of BHB1 stars is shown in blue line and histogram of BHB2 stars is shown in green line. The lower and upper  $x$ -axes is the distance from the cluster center in parsec and arcminutes, respectively. The  $y$ -axes represents the number of stars present in each distance bin of 5 parsec from the cluster center. The vertical dashed line is the tidal radius of the cluster.

center. However, BHB2 stars show their radial distribution up to 25 pc. We find that the BHB1 stars are having almost double counts than the BHB2 stars up to 10 pc. This indicates that the hotter sources are relatively concentrated at the center than the lower  $T_{\text{eff}}$  BHBs.

## 5. FUV bright blue straggler star

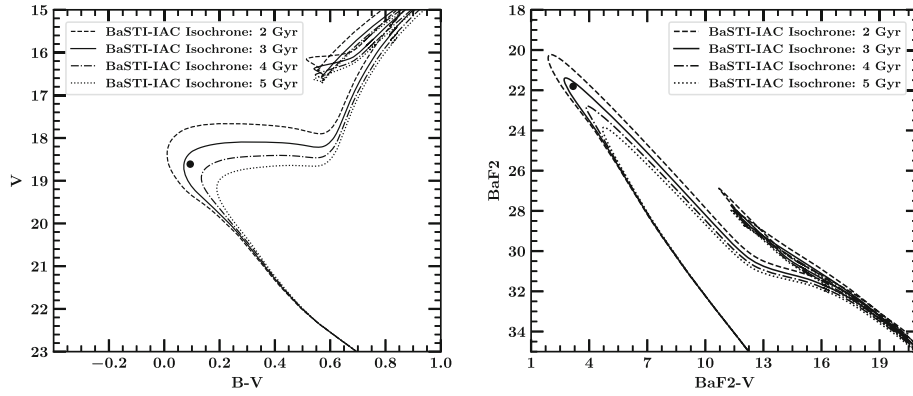
The blue-stragglers (BS) are generally formed in the old age clusters through two formation channels: mass-transfer in binary system in low density environment (Knigge *et al.* 2009; Leigh *et al.* 2013) and merger due to collision between stars in dense environment (Chatterjee *et al.* 2013). They are relatively more massive hydrogen burning MS stars than the normal MS population of the cluster owing to the higher mass and consequently shifted in the upper MS branch in the H-R diagram. The UV photometry is a great tool to identify the BS population in a cluster. The data from UV telescopes (*HST*, *GALEX* and *UVIT*) have been excellent tool to explore BS population and their formation scenario (Ferraro *et al.* 1997, 2003; Dieball *et al.* 2010; Schiavon *et al.* 2012; Gosnell *et al.* 2015; Subramaniam *et al.* 2017; Sahu

*et al.* 2019a; Kumar *et al.* 2020, 2021; Rani *et al.* 2020). In particular, *UVIT* observations of star clusters are able to identify the hot binary companions (WD, HB, EHB, etc.) of BS stars (Subramaniam *et al.* 2016; Sindhu *et al.* 2019; Sahu *et al.* 2019b; Jadhav *et al.* 2019; Singh *et al.* 2020).

Present observational data-set allows us to identify one FUV bright BSS in the outer region of the cluster, at a distance of 2.15' from the center. The respective colors and magnitudes of the BSS in UV-optical and optical CMDs are shown in Fig. 2. The BS are generally more massive than the stars currently evolving at the MS turn-off and are located along the brighter and hotter extension of the MS locus, mimicking the location of intermediate-age stars. Therefore, we generated relatively younger ages BaSTI-IAC isochrones by adopting the same assumptions concerning metallicity and alpha-element enhancement as the old stellar component. The comparison between model predictions and the empirical data in the  $B-V$  vs.  $V$  and  $\text{BaF}_2-V$  vs.  $\text{BaF}_2$  CMDs is shown in Fig. 5. We see that the isochrone with age 3 Gyr is lying closer to the observed BSS. Based on the observed location of the BSS in optical and UV-optical CMDs, we extracted the mass, luminosity, effective temperature of the BSS from the 3 Gyr BaSTI-IAC isochrone which are  $M = 1.14 \pm 0.04 M_{\odot}$ ,  $\log L/L_{\odot} = 0.9 \pm 0.1$ , and  $\log T_{\text{eff}}/K = 3.93 \pm 0.01$ , respectively.

We obtained optical photometry of the FUV bright BSS from various archival catalogs and performed spectral energy distribution (SED) fitting on the observed fluxes to the model generated fluxes for different filters from UV to near-IR bandwidth. We used an online SED fitting tool, VO SED analyzer (VOSA, Bayo *et al.* 2008), to perform the SED fit. It uses the chi-square minimisation technique to fit the observed fluxes on the model fluxes generated from the theoretical spectra incorporating various stellar atmosphere models. We used the Kurucz stellar atmosphere model, ATLAS9 (Kurucz model, Castelli & Kurucz 2003) grids (spectra) with model parameters in the following range,  $T_{\text{eff}}$ : 3500 K to 50000 K,  $[\text{Fe}/\text{H}]$ :  $-2.0$  and  $-1.5$  (nearest to the cluster metallicity), and  $\log(g)$ : 0.0 to 5.0, respectively. The observed fluxes of different filters used in the SED fit is given in Table 2. We have used observed fluxes from 23 filters in the wavelength range 1800 Å to 11800 Å to perform the SED fit.

In Fig. 6, we have shown the best-fitted Kurucz model spectra of 8500 K on the observed fluxes in gray line. The model fluxes calculated for various filters using the Kurucz model spectra are shown in

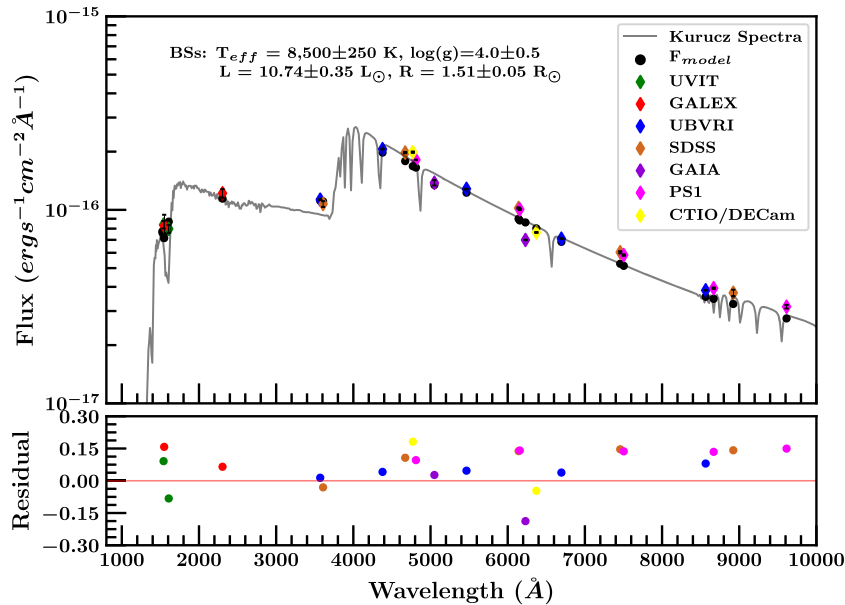


**Figure 5.** Optical and UV-optical CMDs are shown in left and right panels, respectively. The FUV bright BSS of NGC 4147 is shown in the black solid in both CMDs. The BaSTI-IAC isochrones with the same parameters as mentioned in Fig. 2 but with younger ages, i.e. 2 Gyr (dashed line), 3 Gyr (solid line), 4 Gyr (dash-dotted line), and 5 Gyr (dotted line), are over-plotted in both CMDs.

**Table 2.** List of the telescopes and their filters used in the SED fit.

Telescope	Filters	Wavelength range in Å	Reference
UVIT/AstroSat	BaF <sub>2</sub> , sapphire	1350–1800	This paper
<i>GALEX</i>	FUV, NUV	1350–3000	Schiavon <i>et al.</i> (2012)
CFHT-3.6m	U, B, V, R, I	3000–11800	Stetson <i>et al.</i> (2019)
SDSS	u, g, r, i, z	3000–10800	DR9, Ahn <i>et al.</i> (2012)
GAIA	G, BP	3300–10600	Gaia Collaboration <i>et al.</i> (2018)
PAN-STARRS	g, r, i, z, y	3900–10800	Chambers <i>et al.</i> (2016)
CTIO/DECam	g, r	3925–7233	DECam Legacy Survey DR3 <sup>a</sup>

<sup>a</sup><https://www.legacysurvey.org/dr3/description/>



**Figure 6.** The SED fit of BSS using the Kurucz model is shown in the upper panel. The bottom panel shows the residue (deviation) of the observed fluxes from the Kurucz model fluxes.

**Table 3.** Various physical parameters of the BSS derived from the SED fit and BaSTI-IAC isochrones are listed in the table.

ID	RAJ2000 (degree)	DEJ2000 (degree)	Fit	$T_{\text{eff}}$ (K)	Luminosity ( $L_{\odot}$ )	Radius ( $R_{\odot}$ )	$\log(g)$ (dex)	Mass ( $M_{\odot}$ )
BSS01	182.5633	18.5480	Isochrone	$8511 \pm 196$	$7.94 \pm 1.83$	$1.30 \pm 0.07$	–	$1.14 \pm 0.04$
			SED	$8500 \pm 250$	$10.74 \pm 0.35$	$1.51 \pm 0.05$	$4.0 \pm 0.5$	$1.23 \pm 0.02$

**Table 4.** UVIT photometry table of the BHB1, BHB2 and FUV bright BSS containing their respective positions, the extinction-corrected magnitudes and the magnitude errors.

ID	RAJ2000	DEJ2000	BaF <sub>2</sub>	eBaF <sub>2</sub>	Sapphire	eSapphire	Silica	eSilica
HB01	182.5358	18.5215	19.57	0.01	19.37	0.02	19.35	0.02
HB02	182.5117	18.5330	19.79	0.02	19.62	0.02	19.44	0.02
HB03	182.5076	18.5353	19.83	0.02	19.55	0.02	19.41	0.02
HB04	182.5239	18.5356	19.43	0.01	19.23	0.02	19.55	0.03
HB05	182.5231	18.5393	19.46	0.01	19.13	0.01	19.31	0.03
HB06	182.5243	18.5396	19.62	0.01	19.48	0.02	19.29	0.03
HB07	182.5302	18.5413	19.45	0.01	19.27	0.01	19.45	0.02
HB08	182.5229	18.5422	19.53	0.01	19.22	0.01	19.23	0.03
HB09	182.5309	18.5458	20.05	0.01	19.84	0.02	20.08	0.03
HB10	182.5159	18.5455	19.41	0.02	19.13	0.02	19.15	0.02
HB11	182.5288	18.5465	19.70	0.01	19.47	0.02	19.31	0.03
HB12	182.5297	18.5473	19.97	0.01	19.59	0.02	19.70	0.03
HB13	182.5111	18.5564	19.95	0.01	20.07	0.02	19.73	0.02
HB14	182.5240	18.5567	19.85	0.02	19.40	0.01	19.84	0.03
HB15	182.5394	18.5611	19.74	0.01	19.65	0.01	19.41	0.03
HB16	182.5644	18.5639	19.49	0.01	19.42	0.02	19.44	0.02
HB17	182.5427	18.6696	19.52	0.01	19.23	0.02	19.16	0.02
HB18	182.5103	18.5344	19.54	0.02	19.26	0.01	19.33	0.03
HB19	182.5291	18.5376	19.22	0.03	19.09	0.03	18.98	0.03
HB20	182.5228	18.5379	19.51	0.01	19.22	0.01	19.09	0.03
HB21	182.5315	18.5503	19.49	0.01	19.43	0.02	19.36	0.02
HB22	182.5528	18.5609	19.36	0.01	19.11	0.01	19.33	0.02
HB23	182.5381	18.5407	19.28	0.01	19.14	0.01	19.21	0.03
HB24	182.5279	18.5506	19.44	0.02	19.23	0.04	19.38	0.02
HB25	182.5417	18.5302	20.63	0.03	20.54	0.03	20.05	0.03
HB26	182.5150	18.5352	21.78	0.04	21.13	0.03	20.53	0.03
HB27	182.5268	18.5388	20.77	0.02	20.60	0.03	20.26	0.04
HB28	182.5312	18.5420	21.26	0.03	20.84	0.03	20.82	0.06
HB29	182.5171	18.5441	21.84	0.04	21.21	0.04	20.26	0.04
HB30	182.5097	18.5476	20.59	0.02	20.15	0.02	19.89	0.05
HB31	182.5226	18.5489	20.75	0.03	20.13	0.02	19.91	0.02
HB32	182.4897	18.5769	21.73	0.04	21.17	0.04	20.27	0.04
HB33	182.5668	18.5771	21.60	0.03	21.29	0.03	20.59	0.04
HB34	182.5318	18.5426	20.99	0.03	21.05	0.04	20.07	0.03
HB35	182.5307	18.5430	21.96	0.04	21.55	0.05	–	–
HB36	182.5292	18.5481	21.69	0.04	21.45	0.06	–	–
HB37	182.5123	18.6111	21.48	0.05	21.68	0.05	–	–
BSS01	182.5633	18.5480	21.80	0.05	21.81	0.05	–	–

IDs HB01 to HB24 belong to the BHB1 group and IDs HB25 to HB37 belong to the BHB2 group. The BSS is listed at the end of the table with ID BSS01.



solid black circles and the observed fluxes from the UVIT, *GALEX*, CFHT-3.6 m, SDSS, GAIA, PAN-STARRS, and CTIO/DECam telescopes are shown in green, red, blue, brown, violet, magenta, and yellow diamonds, respectively. In the lower panel of Fig. 6, we have shown the fractional deviation of the observed flux from the model flux and indicated as residual  $[(F_{\text{obs.}} - F_{\text{mod}})/F_{\text{mod}}]$ . The best-fitted Kurucz model spectra with physical parameters  $T_{\text{eff}} = 8500$  K,  $[\text{Fe}/\text{H}] = -2.0$  and  $\log(g) = 4.0$ , was obtained with a reduced chi-square value of 9.3. The deviation of the observed fluxes from the theoretical fluxes for most of the filters is less than 15%. But for a few filters (e.g., CTIO/DECam  $g$  and Gaia  $G$  filters) the deviation is more than or around 15% which might be the reason for the high value of  $\chi^2$ . However, we can see that the overall photometric magnitudes are well within the maximum photometric errors of 20% (0.2 mag). We extracted bolometric luminosity and radius of the FUV bright BSS from the slope of the best fitted spectrum using scaling relation  $(R/d)^2$  used in the SED fit, where  $R$  is the radius of the source and  $d$  is the distance. The distance of the cluster = 19.0 kpc and extinction value  $A_v = 0.068$  magnitude were used in the SED fit. We have also estimated the mass of the BSS using the  $T_{\text{eff}}$ , luminosity and radius derived from the SED analysis. The derived physical parameters from the SED fit and the BaSTI-IAC isochrone fitting are listed in Table 3. We find that the  $T_{\text{eff}}$ , mass and luminosity of the BSS, derived from the SED fit and BaSTI-IAC isochrone fitting on CMDs are matching well within the error range.

## 6. Summary and conclusion

UV analysis of the GGC NGC 4147 is presented using observations in three FUV filters of UVIT. A catalog of the UVIT detected BHBs, and BSS along with their FUV magnitudes in three UVIT filters is provided in Table 4. It is found that all the FUV bright BHBs belong to the 2G population as defined by Villanova *et al.* (2016) and they lie in two sub-populations on the UV-optical CMDs with a count ratio of 24:13 (BHB1:BHB2). The derived ranges of  $T_{\text{eff}}$  of BHB1 and BHB2 are 10000–11300 K and 8000–9500 K, respectively with a gap of 500 K between both the sub-populations. Spectroscopic analysis of the sources is required to investigate further on the distinction between the two sub-population of BHBs. The physical parameters of the FUV bright BSS were derived

by fitting younger BaSTI-IAC isochrones in optical and UV-optical CMDs, and SED fitting of the observed fluxes from UV to near-IR wavelengths with the Kurucz model. The derived physical parameters of the FUV bright BSS are listed in Table 3.

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## References

- Ahn C. P., Alexandroff R., Allende Prieto C. *et al.* 2012, ApJS, 203, 21
- Arellano Ferro A., Arévalo M. J., Lázaro C. *et al.* 2004, Revista Mexicana de Astronomía y Astrofísica, 40, 209
- Arellano Ferro A., Rojas Galindo F. C., Muneer S., Giridhar S. 2018, Revista Mexicana de Astronomía y Astrofísica, 54, 325
- Auriere M., Lauzeral C. 1991, A&A, 244, 303
- Bastian N., Lardo C. 2018, ARA&A, 56, 83
- Bayo A., Rodrigo C., Barrado Y Navascués D. *et al.* 2008, A&A, 492, 277
- Behr B. B., Djorgovski S. G., Cohen J. G. *et al.* 2000, ApJ, 528, 849
- Brown T. M., Cassisi S., D'Antona F. *et al.* 2016, ApJ, 822, 44
- Cardelli J. A., Clayton G. C., Mathis J. S. 1989, ApJ, 345, 245
- Cassisi S., Salaris M. 2020, A&AR, 28, 5

- Castelli F., Kurucz R. L. 2003, in Piskunov N., Weiss W. W., Gray D. F., eds, *Modelling of Stellar Atmospheres*, Vol. 210, A20
- Catelan M., Borissova J., Sweigart A. V., Spassova N. 1998, *ApJ*, 494, 265
- Chambers K. C., Magnier E. A., Metcalfe N. *et al.* 2016, arXiv e-prints, [arXiv:1612.05560](https://arxiv.org/abs/1612.05560)
- Chatterjee S., Rasio F. A., Sills A., Glebbeek E. 2013, *ApJ*, 777, 106
- Clement C. 2017, in Catelan M., Gieren W., eds, *European Physical Journal Web of Conferences*, Vol. 152, *Wide-Field Variability Surveys: A 21st Century Perspective*, EDP Sciences, France, 01021
- Dieball A., Long K. S., Knigge C., Thomson G. S., Zurek D. R. 2010, *ApJ*, 710, 332
- Ferraro F. R., Sills A., Rood R. T., Paltrinieri B., Buonanno R. 2003, *ApJ*, 588, 464
- Ferraro F. R., Paltrinieri B., Fusi Pecci F. *et al.* 1997, *A&A*, 324, 915
- Gaia Collaboration, Brown A. G. A., Vallenari A. *et al.* 2018, *A&A*, 616, A1
- Gosnell N. M., Mathieu R. D., Geller A. M. *et al.* 2015, *ApJ*, 814, 163
- Gratton R., Bragaglia A., Carretta E. *et al.* 2019, *A&AR*, 27, 8
- Harris W. E. 2010, arXiv e-prints, [arXiv:1012.3224](https://arxiv.org/abs/1012.3224)
- Hidalgo S. L., Pietrinferni A., Cassisi S. *et al.* 2018, *ApJ*, 856, 125
- Ivans I. I. 2009, in *American Astronomical Society Meeting Abstracts*, Vol. 213, *American Astronomical Society Meeting Abstracts #213*, 341.06
- Jadhav V. V., Sindhu N., Subramaniam A. 2019, *ApJ*, 886, 13
- Jain R., Vig S., Ghosh S. K. 2019, *MNRAS*, 485, 2877
- Knigge C., Leigh N., Sills A. 2009, *Nature*, 457, 288
- Kumar A. *et al.* 2012, in Takahashi T., Murray S. S., den Herder J.-W. A., eds, *Proc. SPIE Conf. Ser.*, Vol. 8443, *Space Telescopes and Instrumentation 2012: Ultraviolet to Gamma Ray*, SPIE, Bellingham, p. 84431N
- Kumar R., Pradhan A. C., Parthasarathy M., Ojha D. K., Mohapatra A., Murthy J. 2020, in Bragaglia A., Davies M., Sills A., Vesperini E., eds, *Proc. IAU Symp. 351, Star Clusters: From the Milky Way to the Early Universe*, Cambridge University Press, Cambridge, UK, p. 464
- Kumar R., Pradhan A. C., Mohapatra A. *et al.* 2021, *MNRAS* 502, 313
- Lata S., Pandey A. K., Pandey J. C. *et al.* 2019, *AJ*, 158, 51
- Leigh N., Knigge C., Sills A. *et al.* 2013, *MNRAS*, 428, 897
- Marino A. F., Milone A. P., Renzini A. *et al.* 2019, *MNRAS*, 487, 3815
- Martell S. L., Smith G. H., Briley M. M. 2008, *AJ*, 136, 2522
- Milone A. P., Piotto G., Renzini A. *et al.* 2017, *MNRAS*, 464, 3636
- Nardiello D., Libralato M., Piotto G. *et al.* 2018, *MNRAS*, 481, 3382
- Pietrinferni A., Hidalgo S., Cassisi S. *et al.* 2021, *ApJ*, in press, arXiv e-print, [arXiv:2012.10085](https://arxiv.org/abs/2012.10085)
- Piotto G., Zoccali M., King I. R. *et al.* 1999, *AJ*, 118, 1727
- Postma J. E., Leahy D. 2017, *PASP*, 129, 115002
- Rani S., Pandey G., Subramaniam A., Sahu S., Kameswara Rao N. 2020, *J. Astrophys. Astr.*, 41, 35
- Sahu S., Subramaniam A., Côté P., Rao N. K., Stetson P. B. 2019a, *MNRAS*, 482, 1080
- Sahu S., Subramaniam A., Simunovic M. *et al.* 2019b, *ApJ*, 876, 34
- Schiavon R. P., Dalessandro E., Sohn S. T. *et al.* 2012, *AJ*, 143, 121
- Schlafly E. F., Finkbeiner D. P. 2011, *ApJ*, 737, 103
- Sindhu N., Subramaniam A., Jadhav V. V. *et al.* 2019, *ApJ*, 882, 43
- Singh G., Sahu S., Subramaniam A., Yadav R. K. S. 2020, *ApJ*, 905, 44
- Stetson P. B. 1987, *PASP*, 99, 191
- Stetson P. B., Catelan M., Smith H. A. 2005, *PASP*, 117, 1325
- Stetson P. B., Pancino E., Zocchi A., Sanna N., Monelli M. 2019, *MNRAS*, 485, 3042
- Subramaniam A., Sindhu N., Tandon S. N. *et al.* 2016, *ApJL*, 833, L27
- Subramaniam A., Sahu S., Postma J. E. *et al.* 2017, *AJ*, 154, 233
- Suntzeff N. B., Kraft R. P., Kinman T. D. 1988, *AJ*, 95, 91
- Tailo M., Milone A. P., Lagioia E. P. *et al.* 2020, *MNRAS*, 498, 5745
- Tandon S. N., Subramaniam A., Girish V. *et al.* 2017, *AJ*, 154, 128
- Tandon S. N., Postma J., Joseph P. *et al.* 2020, *AJ*, 159, 158
- Villanova S., Monaco L., Moni Bidin C., Assmann P. 2016, *MNRAS*, 460, 2351