

## Line Shape Variability in a Sample of AGN with Broad Lines

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Received 4 July 2015; accepted 1 October 2015

DOI: 10.1007/s12036-015-9360-y

**Abstract.** The spectral variability of active galactic nuclei (AGN) is one of the key features that enables us to study in more detail, the structure of AGN emitting regions. Especially, the broad line profiles that vary both in flux and shape, give us invaluable information about the kinematics and geometry of the broad line region (BLR) where these lines are originating from. We give here a comparative review of the line shape variability in a sample of five type 1 AGNs, those with broad emission lines in their spectra, of the data obtained from the international long-term optical monitoring campaign coordinated by the Special Astrophysical Observatory of the Russian Academy of Science. The main aim of this campaign is to study the physics and kinematics of the BLR on a uniform data set, focusing on the problems of the photoionization heating of the BLR and its geometry, where, in this paper, we give for a first time, a comparative analysis of the variability of five type 1 AGNs, discussing their complex BLR physics and geometry in the framework of the estimates of the supermassive black hole mass in AGN.

*Key words.* Galaxies: active-galaxies—quasar: individual (Arp 102B, 3C 390.3, NGC 5548, NGC 4151, Ark 564)—line: profiles.

### 1. Introduction

In spite of decades of intensive investigations, the broad line region (BLR) of active galactic nuclei (AGN) is yet not fully understood. The direct detection of the BLR

remains a challenge for modern instruments, since the angular size of the BLR is less than 0.001 arcsec even for the closest AGNs. The only information from the BLR comes in the form of the broad emission lines (BEL), that are a very prominent feature in the spectra of the so-called type 1 AGN (Osterbrock & Ferland 2006). Nevertheless, the investigations of the BEL's properties (flux and shape), and especially of their variability, indicates that the BLR is linked with the accretion process onto the supermassive black hole in the center of an AGN. First of all, the BLR gas is photoionized by the continuum radiation from the accretion disk, and secondly, at least part of the BLR follows its geometry, i.e. even though the BLR geometry is not ubiquitously described it is believed that one part of the BLR is a part of the accretion disk or has a disk-like geometry (Popović *et al.* 2004).

An important method to indirectly map the geometry and kinematics of the the BLR is the 'reverberation mapping' (Blandford & McKee 1982; Gaskell 1988; Peterson 1993; Peterson *et al.* 2004), a method that is based on multiple spectroscopic observation and the variability of spectral properties. In particular, the BEL's flux varies with respect to the ionizing continuum flux with a certain time delay (of the order of days to weeks), due to the light-travel time from the source of the continuum photon to the source of the BEL photon, that is the BLR. Therefore, by applying the cross-correlation function (CCF) between the continuum and BEL light curves an estimate of the time-lag  $\tau$  between the two signals can be obtained, which is basically the photometric BLR radius  $R_{\text{BLR}} = c\tau$ , where  $c$  is the speed of light. This is a powerful tool to observe the unobservable, as one can estimate the size of the BLR and consequently the mass of the supermassive black hole, since the BLR gas is assumed to be virialized, and the method has been exploited in many papers (see e.g. Kaspi *et al.* 2000; Peterson *et al.* 2004; Bentz *et al.* 2009; Doroshenko *et al.* 2012; Kollatschny *et al.* 2014). The reverberation-mapped AGNs are of particular importance because they anchor the scaling relationships (BLR radius-luminosity relationship, see e.g. Kaspi *et al.* 2000; Bentz *et al.* 2006) that allow the estimate of the BLR radius, and the supermassive black hole mass, from a single-epoch spectrum.

However, this method is based on several assumptions, which makes it still pretty uncertain, especially for the supermassive black hole mass estimate. It is assumed that the BLR gas is photoionized by the continuum coming from the central source, and that it is gravitationally bound to the supermassive black hole so that the virial theorem can be used (Gaskell 1988). Moreover, it strongly depends on the geometry of the BLR, and the kinematics and structure of the BLR is not ubiquitously described (see e.g. Sulentic *et al.* 2000; Gaskell 2009 and reference therein). The BLR is probably very complex, often with evidence for multiple components (e.g. Sulentic *et al.* 2000; Ilić *et al.* 2006; Popović *et al.* 2004; Bon *et al.* 2009), and can have systematic motions such as infalls, outflows, circular motions, which should be detected in the BEL profiles. If the line profile is varying, the change of the line shape should be according to the change in the BLR geometry and kinematics. In principle, the reverberation assumption can be tested, e.g. one test is if the BEL fluxes correlate with the continuum flux. Therefore, detailed studies of different broad line profiles is required, and this can be done with long-term (of the order of a decade and longer) monitoring of type 1 AGNs with different BEL properties.

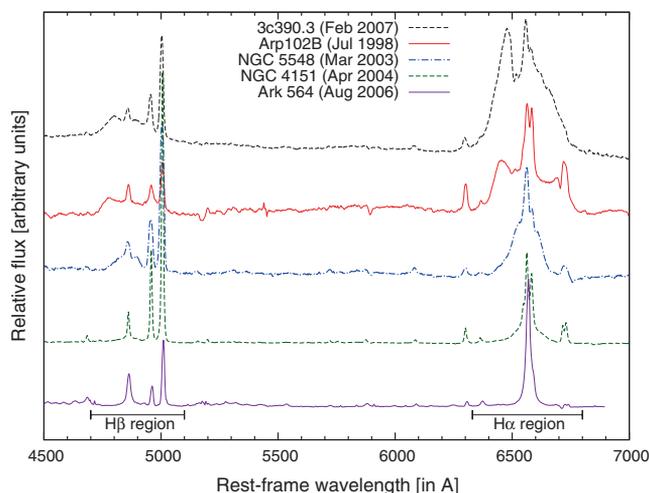
In this paper, we give a comparative overview of our investigation based on the international long-term optical monitoring campaign, coordinated by the Special

Astrophysical Observatory of the Russian Academy of Science (SAO RAS). Within this campaign, several type 1 AGN have been monitored for decades, out of which data were published for 5 objects (Shapovalova *et al.* 2001, 2004, 2008, 2010a, 2012, 2013). The objects have been selected so that representatives of type 1 AGNs with different BEL profiles are in the sample: Arp 102B and 3C 390.3 have BELs with double-peaked line profiles, Ark 564 is a Narrow-Line Seyfert 1 (NLSy1), and NGC 4151 and NGC 5548 are well-studied Seyfert 1.5 with strongly variable BEL profiles, with bumps and asymmetries present in the line profiles (see Figure 1). All 5 AGNs have been discussed individually in referenced papers, but here we aim to give for the first time a comparative analysis of their variability properties, that is: (i) to summarize the already obtained most important results of their BEL variability, (ii) to give new comparative analysis of the BEL properties, (iii) to discuss their complex BLR physics and geometry, all in the framework of the estimations of the mass of the supermassive black hole in AGN, and suggestions for future exploitation of our large data sets. Our data sets are observed, reduced and analyzed using the same procedures, thus the total analyzed data form an uniform set of data.

The paper is organized as follows: in section 2, we briefly report on observation and data reduction, in section 3, we present the main result, which are discussed in section 4, and in section 5, we outline our conclusions and future investigations.

## 2. Observations and data sets

With the international long-term optical monitoring campaign, coordinated by SAO RAS several type 1 AGN have been monitored for decades, out of which data were published for 5 objects in 2001–2013 (see Table 1 for details). The high quality spectra were taken with 6 different telescopes: (i) 6-m and 1-m telescopes of SAO RAS; (ii) 2.1-m telescope of Guillermo Haro Astrophysical Observatory, Mexico;



**Figure 1.** Comparison of optical spectra for all 5 objects (name and month of the observation are denoted), given in arbitrary flux units and shifted for better display. Rest-frame wavelengths are displayed on the X-axis, and the spectral region of H $\beta$  and H $\alpha$  line are marked.

**Table 1.** Spectral characteristics and the variability of the observed objects: object name and the monitoring period, redshift, AGN type, line profile characteristics and the FWHM of the mean spectrum, photometric BLR radius  $R_{\text{BLR}}^{\text{ph}}$ , maximal and minimal flux ratio  $F_{\text{max}}/F_{\text{min}}$ , the variability parameter  $F_{\text{var}}$ , mean continuum luminosity at 5100 Å, and the main references. Available values for both H $\beta$  and H $\alpha$  line are listed.

Object (period/years)	$z$	AGN type	Line profile shape FWHM (km/s)	$c\tau_{\text{BLR}}$ [ld] H $\beta$ /H $\alpha$	$F_{\text{max}}/F_{\text{min}}$ H $\beta$ /H $\alpha$	$F_{\text{var}}$ H $\beta$ /H $\alpha$	$\lambda L_{\lambda}(5100)$ ( $10^{44}$ erg/s)	Ref.
NGC 5548 (1996-2002)	0.0172	Sy 1.0-1.8	Strong shoulders 6300	$49^{+19}_{-8}/27^{+14}_{-6}$	$4.9 \pm 0.3/3.5 \pm 0.3$	$0.33/0.30$	$0.40 \pm 0.12$	Shapovalova et al. (2004) Kovačević et al. (2014)
NGC 4151 (1996-2006)	0.0033	Sy 1.5-1.8	Absorption, bumps 6110/4650	$5^{+28}_{-5}/6^{+27}_{-6}$	$4.8 \pm 0.2/3.1 \pm 0.3$	$0.42/0.31$	$0.05 \pm 0.03$	Shapovalova et al. (2008)
3C390.3 (1995-2007)	0.0561	Radio-loud	Double-peaked 11900/11000	$96^{+28}_{-47}/120^{+18}_{-18}$	$4.7 \pm 0.4/3.4 \pm 0.6$	$0.38/0.35$	$0.90 \pm 0.42$	Shapovalova et al. (2010a)
Ark 564 (1999-2010)	0.0247	NLSy1	Strong FeII 960/800	$4^{+27}_{-4}/5^{+6}_{-14}$	$1.6 \pm 0.1/1.5 \pm 0.1$	$0.07/0.08$	$0.36 \pm 0.04$	Shapovalova et al. (2012)
Arp 102B (1987-2013)	0.0242	LINER	Double-peaked 15900/14300	$15^{+20}_{-15}/21^{+14}_{-14}$	$3.0 \pm 0.2/2.4 \pm 0.2$	$0.21/0.20$	$0.11 \pm 0.01$	Shapovalova et al. (2013)

(iii) 2.1-m telescope of the Observatorio Astronómico Nacional at San Pedro Martir, Mexico, and (iv) the 3.5-m and 2.2-m telescopes of Calar Alto Observatory, Spain. The data acquisition, reduction and different calibration (e.g. corrections for different position angle, seeing and aperture effects), and flux measurements were described in detail in Shapovalova *et al.* (2001, 2004, 2008, 2010a, 2012, 2013). Table 1 gives a summary of the data (object name, monitored period, redshift and type, line properties and references).

For all objects we measured different line parameters, such as total line fluxes, broad line fluxes,<sup>1</sup> line widths, and line-flux time-lags, that will be investigated here (see references listed in Table 1 for details on other line parameters). Some of the measured H $\beta$  and H $\alpha$  line parameters are given in Table 1: the full-width half-maximum (FWHM) of the mean spectrum, the photometric radius of the BLR  $R_{\text{BLR}} = c\tau_{\text{BLR}}$ , where  $\tau_{\text{BLR}}$  is the line-flux time-lag obtained from the CCF analysis, the maximal-to-minimal flux ratio  $F_{\text{max}}/F_{\text{min}}$ , the variability parameter  $F_{\text{var}}$ , and the continuum luminosity at 5100 Å. The variability parameter  $F_{\text{var}}$  is used to estimate the amount of the line-variability and is calculated according to O’Brien *et al.* (1998). For the estimates of time-lags from the CCF analysis a  $Z$ -transformed Discrete Correlation Function (ZDCF) was used (Alexander 1997). Luminosities are calculated using the mean continuum flux at 5100 Å and the online calculator for luminosity distance (Wright 2006), for which we adopted the cosmological parameters  $\Omega = 0.286$ ,  $\Omega_{\Lambda} = 0.714$  and  $\Omega_k = 0$ , and a Hubble constant,  $H_0 = 69.6 \text{ km/s Mpc}^{-1}$ . All flux measurements from this campaign are publicly available in publications listed in Table 1 and in the corresponding VizieR Online Data Catalog.

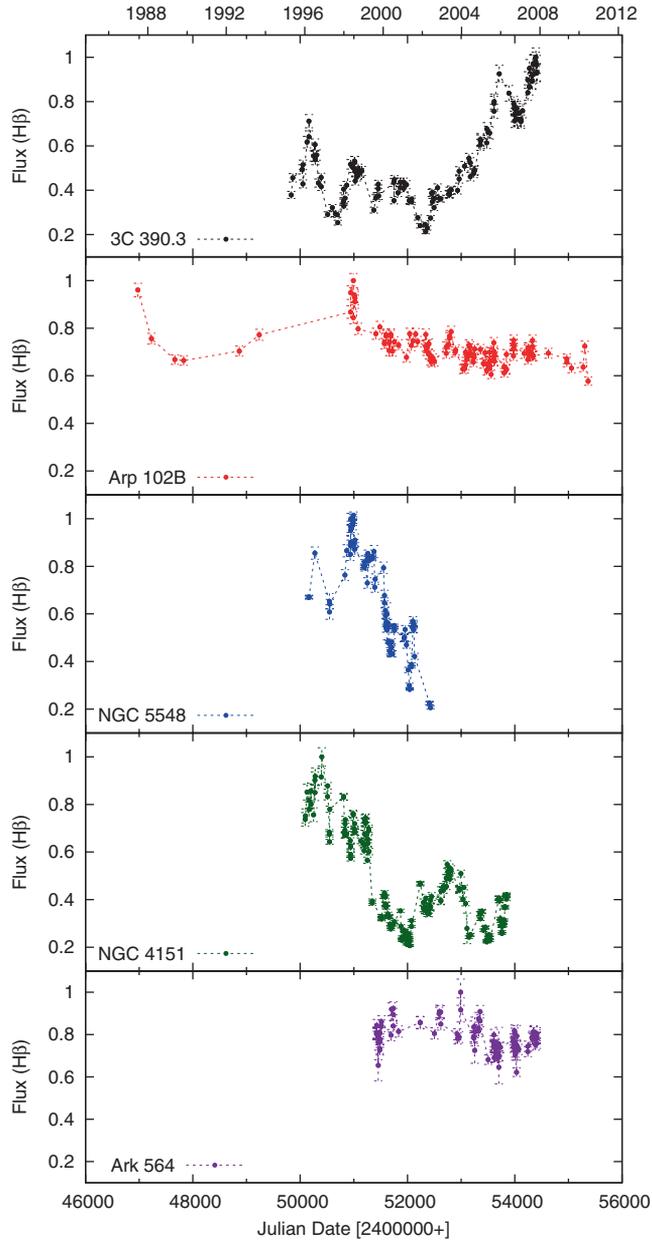
The same procedures were performed for the observations, data reduction, measurements and analysis, so that the resulting data sets of these 5 objects basically are uniform. This makes them valid for further analysis and comparison.

### 3. Results

Our sample consists of 5 type 1 AGN with different BEL profiles (Figure 1): Arp 102B and 3C 390.3 have BELs with double-peaked line profiles, Ark 564 is a NLSy1, and NGC 4151 and NGC 5548 are strongly variable Seyfert 1.5. The results of our long-term monitoring campaign are summarized in Table 1 and Figures 1 and 2, that show the comparison of optical spectra for all 5 objects, given in arbitrary flux units and shifted for better display (Fig. 1), and the H $\beta$  line flux light curves for all 5 objects (Fig. 2), where the line flux and corresponding errors are normalized to the maximal flux for better comparison.

As in Kollatschny *et al.* (2006), in Figure 3, we plot the ratio of the maximal-to-minimal line flux against the FWHM of H $\beta$  (circles) and H $\alpha$  line (diamonds). We divided the  $F_{\text{max}}/F_{\text{min}}$  vs. FWHM plane according to the line FWHM in three groups of objects: (1) NLSy1 having FWHM < 2000 km/s (Osterbrock 1987), (2) AGN with broad emission lines (Seyfert 1) with 2000 < FWHM < 10000 km/s, and (3) double-peaked line AGN (DPL AGN) with FWHM > 10000 km/s. Figure 3 gives an

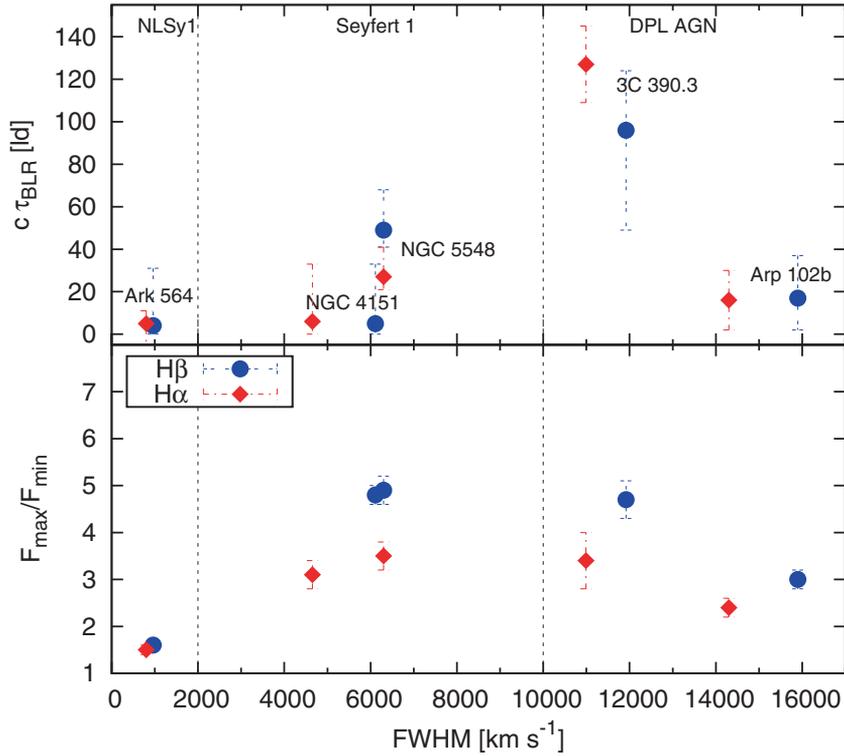
<sup>1</sup>To study the broad components showing the main BLR characteristics, the narrow lines and the forbidden lines were subtracted either by constructing observational templates or by multi-Gaussian fittings.



**Figure 2.** H $\beta$  line light curves for all 5 objects (3C 390.3, Arp 102B, NGC 5548, NGC 4151 and Ark 564). The line flux and corresponding error-bars are normalized to the maximal flux for better comparison and X-axis shows time in units of modified Julian date.

indication that there are different types of objects (with different geometry and type of variability) that should be considered simultaneously.

In order to test the photoionization assumption, we plot in Figure 4 the correlation between the H $\beta$  line and continuum flux for all objects (name indicated on

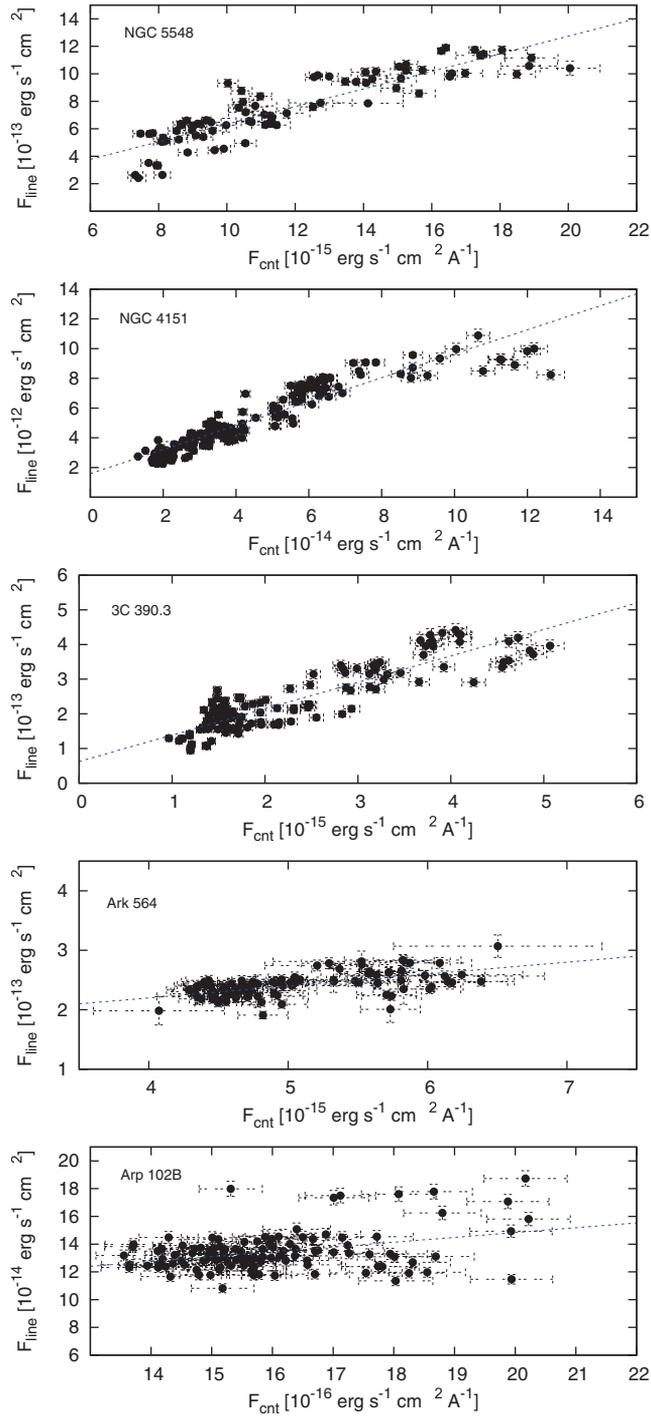


**Figure 3.** Photometric radius of the BLR  $c\tau_{\text{BLR}}$  (*upper panel*) and the variability indicator, the ratio of maximal-to-minimal flux  $F_{\text{max}}/F_{\text{min}}$  (*bottom panel*) versus the FWHM of H $\beta$  (circles) and H $\alpha$  (diamonds) lines for all 5 objects. The vertical dashed lines correspond to the line FWHM of 2000 and 10000 km/s.

each plot). The correlation coefficient and the corresponding P-value are listed in Table 2. The two objects, Arp 102B and Ark 564 have very low correlation coefficients (0.37 and 0.59, respectively), indicating a weak response of the line to the continuum variability.

In principle, the geometry of the BLR should be seen in the line shape, e.g. in the most obvious case if the rotational motion is present in the BLR, then the line profile should have two peaks, the red one corresponding to the gas moving away, and the blue one moving towards the observer. Moreover, if there is a change in the BLR geometry and kinematics, the line shape should vary accordingly. In order to show how the line profiles vary in two DPL objects, we plot the mean and rms profile for the H $\beta$  line of 3C 390.3 (Figure 5, upper panel) and Arp 102B (Figure 5, bottom panel). From the shape and intensity of the rms profiles, it is obvious that in case of 3C 390.3 the line profiles vary significantly, while in Arp 102B, the line profiles remain almost unchanged.

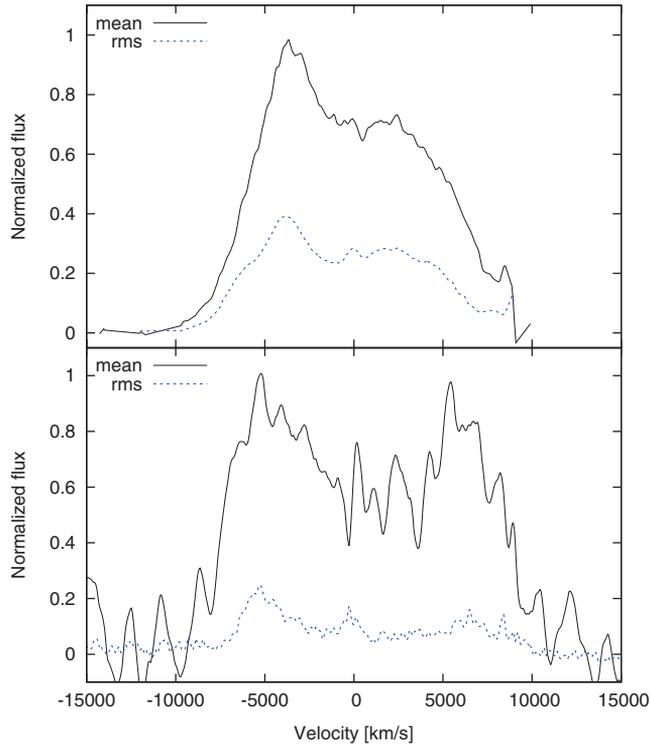
Finally, in Figure 6, the radius of the BLR for H $\beta$  line versus the continuum luminosity at 5100 Å, is given. The solid line represents the scaling relation from Bentz *et al.* (2006),  $\log R_{\text{BLR}} = K + \alpha \log[\lambda L_{\lambda}(5100)/10^{44}]$ , where  $\alpha$  is the slope of the BLR radius-luminosity relationship and  $K$  is the scaling factor, and in Figure 6,



**Figure 4.** Correlation between the line and continuum flux for the  $H\beta$  line for all 5 objects (NGC 5548, NGC 4151, 3C 390.3, Ark 564 and Arp 102B). The dashed line represents the linear best fit through all data.

**Table 2.** Line and continuum flux correlations for all 5 objects. The correlation coefficient and the corresponding  $P_{\text{null}}$  value (in brackets) are given for each pair of data.

Object	H $\beta$ vs. $F_{\text{cnt}}$	H $\alpha$ vs. $F_{\text{cnt}}$
NGC 5548	0.90 (0.1E-28)	0.85 (0.2E-13)
NGC 4151	0.94 (0.0)	0.88 (0.0)
3C390.3	0.91 (1.1E-19)	0.82 (1.1E-12)
Ark 564	0.59 (0.7E-9)	0.71 (0.8E-8)
Arp 102B	0.37 (0.5E-4)	0.25 (0.02)

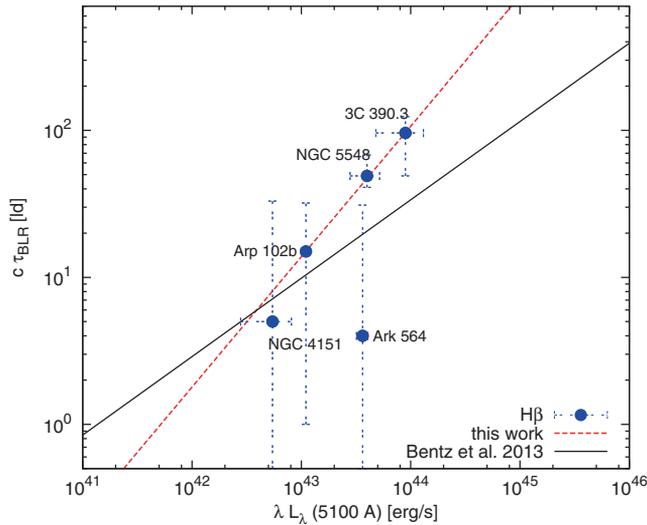


**Figure 5.** Mean (solid line) and rms (dashed line) line profiles for the H $\beta$  line of 3C 390.3 (*upper panel*) and Arp 102B (*bottom panel*).

$\alpha = 0.533$  and  $K = 1.527$  from Bentz *et al.* (2013) are plotted. The dashed line is a simple linear fit of the above equation (not considering the error-bars) through all objects except from Ark 564, obtaining the fitting parameters of  $\alpha = 0.887$  and  $K = 2.027$ .

#### 4. Discussion

We have analyzed the variability of the optical spectra for five type 1 AGNs with different properties of the broad emission lines. The data sets were processed in a uniform way, which makes them valid for further analysis and comparison.



**Figure 6.** The photometric radius of the BLR versus the continuum luminosity at 5100 Å, for H $\beta$  line. The solid line represents the BLR radius-luminosity scaling relation from Bentz *et al.* (2013) where  $R_{\text{BLR}} \sim L^{0.533}$ , while the dashed line is our fit through all objects except Ark 564, where  $R_{\text{BLR}} \sim L^{0.887}$ .

During long-term monitoring, that is, by rule longer than a decade, the line and continuum flux for all objects are varying. Most of the objects are strongly varying, i.e. NGC 5548, NGC 4151 and 3C 390.3, that is best seen from the normalized light curves (Figure 2), but also from the ratio of the maximal-to-minimal flux (Figure 3, Table 1) and the variability parameter  $F_{\text{var}}$  that is  $\sim 40\%$  (Table 1). This is supported with the correlation analysis, as these strongly varying objects show significant correlation between the line and continuum fluxes (Figure 4, Table 2).

On the other hand, the line fluxes of the Ark 564, that is a NLSy1, stay almost constant during the monitored period (only changing by  $\sim 7\text{--}8\%$ ), that was also noticed before for this object (Shemmer *et al.* 2001). It is also interesting that in case of Arp 102B, that is, the other extreme object with the largest line widths (FWHM  $\sim 15000$  km/s), the variability is also low ( $\sim 20\%$ ). The lack of variability in both these cases is supported by the weak correlation between the continuum and line fluxes (Figure 4, Table 2), e.g. for H $\beta$  of Arp 102B, it is  $r = 0.37$  ( $P_{\text{null}} = 0.5\text{E-}04$ ), and even worse for H $\alpha$  of the same object ( $r = 0.25$ ,  $P_{\text{null}} = 0.22$ ). The lack of variability can be due to poor sampling of the data, e.g. in case of Arp 102B, the mean sampling is 40 days (Kovačević *et al.* 2014). However, the lack of correlation between the line and continuum flux also points to additional sources of ionization in the BLR apart from the central AGN continuum source. We should note here that our result for observed low variability in NLS1 Ark 564 is in agreement with a comparative study of NLS1 and broad line Seyfert 1 galaxies (BLS1) variability given by Ai *et al.* (2013). They found that the majority of NLS1-type AGN showed variability on timescales from several days to a few years, but with the variability amplitudes smaller than in the case of BLS1-type AGN. Moreover, there are some results in particular study cases that indicate peculiar BLR in NLS1 AGN, as in NLS1 Mrk 493 (see Popović *et al.* 2009)

where the broad lines have not changed in a period of a couple of years (between two observations), and it is suspected that the weak broad  $H\alpha$  and  $H\beta$  may come partly from the unresolved central BLR, but also may be partly produced by the violent starburst in the circumnuclear ring (clearly resolved on HST image). It seems that NLS1-type AGN have some specific physical properties that have to be considered in monitoring campaigns.

Another peculiarity to note is the fact that the line flux appears to flatten at high continuum flux, that is especially seen in the case of NGC 4151 for which the line flux saturates above the continuum flux of  $7 \times 10^{14} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$  (see thorough analysis and discussion in Shapovalova *et al.* (2008)). Here this effect can be also noticed in the case of NGC 5548 and only weakly in 3C 390.3. This implies that either there is more ionizing to optical flux than expected for a typical AGN spectrum, or that, again, lines are not produced purely by the photoionization from the central continuum source.

The photometric radius of the BLR  $c\tau_{\text{BLR}}$  increases with the width of the BEL's profiles (Figure 3), except for the case of Arp 102B. This is in accordance with the general picture of the central engine, as it appears that in objects with lower accretion rate ( $L/L_{\text{edd}}$ ), i.e. those with larger line widths (Marziani *et al.* 2014), the BLR is larger. Arp 102B though remains a mystery, as here some other geometry, apart from the obvious accretion disk geometry that is usually used to explain the double-peaked line profiles, may be considered (see Popović *et al.* 2014). In favor to this is the fact that during more than two decades, the line profiles of this object have not changed significantly (Figure 5, bottom panel). This is not expected in case of the relatively compact accretion disk ( $\sim 1000$  gravitational radii) that was used to model these very broad double-peaked line profiles (see models in Gezari *et al.* 2007; Popović *et al.* 2014). Additionally, there is an indication that an outflow in the BLR of Arp 102B is detected as shown in Figure 15 of Popović *et al.* (2014).

Finally, the radius–luminosity relationship plotted for these five objects shows that a linear trend is present (Figure 6). Obviously the NLSy1 Ark 564 is clearly an outlier. This is expected due to the low flux variability of this object ( $\sim 7\%$  for  $H\beta$  line) and poor significance of the obtained BLR radius. Therefore, a simple fit was done through all objects except Ark 564, and we obtained  $\alpha = 0.887$  (Figure 6). This result is different from Bentz *et al.* (2013), who obtained  $\alpha = 0.553$ , that is also plotted as a solid line in Figure 6. This illustrates that considering only this very low-luminosity end ( $\lambda L_{\lambda}(5100) \sim 10^{43-44} \text{ erg/s}$ ) of the empirical radius–luminosity scaling relation is very sensitive to the measured continuum luminosity and BLR radius. One reason could be the contribution of the host-galaxy continuum flux, that was not considered here, apart from subtracting the extended source correction factor  $G(g)$ , that is an aperture-dependent correction factor standardly used in our campaign to account for the host-galaxy contribution. Moreover, the sample is far from being statistically significant. However, this still illustrates that new results of the monitoring campaigns may help in obtaining more accurate results and all AGN properties should be taken into account (e.g. level of line variability or line production mechanism), as well as obtaining direct distances needed for calculations of the luminosity as discussed in Bentz *et al.* (2013).

It is obvious that for some objects, the BLR geometry is changing during decades of observation, e.g. in case of 3C 390.3 the rms profile of  $H\beta$  shows strong variability especially in the blue peak (Popović *et al.* 2011). One possible scenario to model this

variability is with the orbiting bright spots in accretion disk (Jovanović *et al.* 2010). Other proposed scenario, e.g. for NGC 4151 that also has strongly variable BEL's profiles (Shapovalova *et al.* 2010b) is that this well-studied object is hosting a binary supermassive black hole with two BLR orbiting around them (Bon *et al.* 2012). Bon *et al.* (2012) found by analyzing the 20-year long NGC 4151 H $\alpha$  light curve and radial velocity curves of the line components, the evidence for a sub-parsec scale supermassive binary black hole system with an orbital period of  $\sim 5800$  days.

## 5. Conclusions

The comparative study of the main results of SAO RAS long-term monitoring campaign of five type 1 AGN with BELs having different properties are given. The variations of the BELs parameter of the H $\beta$  and H $\alpha$  line are investigated with the aim to constrain the properties of the BLR and test the reverberation mapping as a method to estimate the size of the BLR, and consequently the mass of the supermassive black hole in the center of an AGN.

The main conclusions are:

- (i) the five selected objects were varying during the monitored period, where some were strongly variable ( $\sim 40\%$ ) in the spectral line fluxes (e.g. NGC 4151 and 3C 390.3);
- (ii) the BLR is believed to be dominantly photoionized by the central continuum source in the majority of AGN, but in some objects (in this analysis 3 out of 5, i.e. Arp 102B, Ark 564 and NGC 4151), other heating mechanisms such as a shock-ionization due to outflows may contribute to the formation of the BEL;
- (iii) the empirical radius-luminosity scaling relation is very sensitive, especially in this low-luminosity end. It depends on the long-term reverberation monitoring results, so that these investigations continue to be important even though spectroscopy is a very telescope and time-consuming observational technique;
- (iv) during decade(s) of monitored period, the size and geometry of the BLR can change, and this should be considered when BELs are used to estimate the size of the BLR and the mass of the supermassive black hole.

The monitoring results presented here and in several other papers pay attention to the importance of the spectroscopic monitoring campaigns in the investigation of the BLR properties and mass estimates of the supermassive black holes in AGN.

## Acknowledgements

This work was supported by the Ministry of Education, Science and Technological Development of Republic of Serbia through the project Astrophysical Spectroscopy of Extragalactic Objects (176001), INTAS (grant N96-0328), RFBR (grants N97-02-17625 N00-02-16272, N03-02-17123, 06-02-16843, N09-02-01136, 12-02-00857a, 12-02-01237a, N15-02-02101), CONACYT research grants 39560, 54480, and 151494. D.I. has been awarded L'Oréal-UNESCO "For Women in Science" National Fellowship for 2014. We would like to thank the anonymous referees for very useful comments.

## References

- Ai, Y. L., Yuan, W., Zhou, H., Wang, T. G., Dong, X.-B., Wang, J. G., Lu, H. L. 2013, *AJ*, **145**, 90A.
- Alexander, T. 1997, in: *Astronomical Time Series*, edited by D. Maoz, A. Sternberg, E. M. Leibowitz (Dordrecht: Kluwer), 163.
- Bentz, M. C., Peterson, B. M., Pogge, R. W., Vestergaard, M., Onken, C. A. 2006, *ApJ*, **644**, 133.
- Bentz, M. C., Peterson, B. M., Netzer, H., Pogge, R. W., Vestergaard, M. 2009, *ApJ*, **697**, 160.
- Bentz, M. C., Denney, K. D., Grier, C. J. *et al.* 2013, *ApJ*, **767**, 149.
- Blandford, R. D., McKee, C. F. 1982, *ApJ*, **255**, 419.
- Bon, E., Jovanović, P., Marziani, P. *et al.* 2012, *ApJ*, **759**, 118.
- Bon, E., Popović, L. Č., Gavrilović, N., Mura, G. L., Mediavilla, E. 2009, *MNRAS*, **400**, 924.
- Doroshenko, V. T., Sergeev, S. G., Klimanov, S. A., Pronik, V. I., Efimov, Y. S. 2012, *MNRAS*, **426**, 416.
- Gaskell, C. M. 1988, *ApJ*, **325**, 114.
- Gaskell, C. M. 2009, *New Astronomy Reviews*, **53**, 140.
- Gezari, S., Halpern, J. P., Eracleous, M. 2007, *ApJS*, **169**, 167.
- Ilić, D., Popović, L. Č., Bon, E., Mediavilla, E. G., Chavushyan, V. H. 2006, *MNRAS*, **371**, 1610.
- Jovanović, P., Popović, L. Č., Stalevski, M., Shapovalova, A. I. 2010, *ApJ*, **718**, 168.
- Kaspi, S., Smith, P. S., Netzer, H. *et al.* 2000, *ApJ*, **533**, 631.
- Kollatschny, W., Ulbrich, K., Zetzl, M., Kaspi, S., Haas, M. 2014, *A&A*, **566**, A106.
- Kollatschny, W., Zetzl, M., Dietrich, M. 2006, *A&A*, **454**, 459.
- Kovačević, A., Popović, L. Č., Shapovalova, A. I. *et al.* 2014, *Advances in Space Research*, **54**, 1414.
- Marziani, P., Sulentic, J. W., Negrete, C. A. *et al.* 2014, *The Astronomical Review*, **9**, 6.
- O'Brien, P. T., Dietrich, M., Leighly, K. *et al.* 1998, *ApJ*, **509**, 163.
- Osterbrock, D. E. 1987, *Lecture Notes in Physics*, **307**, 1.
- Osterbrock, D. E., Ferland, G. J. 2006, *Astrophysics of gaseous nebulae and active galactic nuclei*, 2nd. ed., CA: University Science Books.
- Peterson, B. M. 1993, *PASP*, **105**, 207.
- Peterson, B. M., Ferrarese, L., Gilbert, K. M. *et al.* 2004, *ApJ*, **613**, 682.
- Popović, L. Č., Mediavilla, E., Bon, E., Ilić, D. 2004, *A&A*, **423**, 909.
- Popović, L. Č., Shapovalova, A. I., Ilić, D. *et al.* 2011, *A&A*, **528**, 130.
- Popović, L. Č., Shapovalova, A. I., Ilić, D. *et al.* 2014, *A&A*, **572**, A66.
- Popović, L. Č., Smirnova, A. A., Kovačević, J., Moiseev, A. V., Afanasiev, V. L. 2009, *AJ*, **137**, 3548P.
- Shapovalova, A. I., Burenkov, A. N., Carrasco, L., Chavushyan, V. H., Doroshenko, V. T., Dumont, A. M. *et al.* 2001, *A&A*, **376**, 775.
- Shapovalova, A. I., Doroshenko, V. T., Bochkarev, N. G. *et al.* 2004, *A&A*, **422**, 925.
- Shapovalova, A. I., Popović, L. Č., Collin, S. *et al.* 2008, *A&A*, **486**, 99S.
- Shapovalova, A. I., Popović, L. Č., Burenkov, A. N. *et al.* 2010a, *A&A*, **517**, A42.
- Shapovalova, A. I., Popović, L. Č., Burenkov, A. N. *et al.* 2010b, *A&A*, **509**, A106.
- Shapovalova, A. I., Popović, L. Č., Burenkov, A. N. *et al.* 2012, *ApJS*, **202**, 10.
- Shapovalova, A. I., Popović, L. Č., Burenkov, A. N. *et al.* 2013, *A&A*, **559**, A10.
- Shemmer, O., Romano, P., Bertram, R. *et al.* 2001, *ApJ*, **561**, 162.
- Sulentic, J. W., Marziani, P., Dultzin-Hacyan, D. 2000, *ARAA*, **38**, 521.
- Wright, E. L. 2006, *PASP*, **118**, 1711.